

## Tests of local position invariance using continuously running atomic clocks

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(Received 11 December 2012; published 22 January 2013)

Tests of local position invariance (LPI) made by comparing the relative redshift of atomic clocks based on different atoms have been carried out for a variety of pairs of atomic species. In most cases, several absolute frequency measurements per year are used to look for an annual signal, resulting in tests that can span on order of a decade. By using the output of continuously running clocks, we carry out LPI tests with comparable or higher precision after less than 1.5 years. These include new measurements of the difference in redshift anomalies  $\beta$  for hyperfine transitions in  $^{87}\text{Rb}$  and  $^{133}\text{Cs}$  and in  $^1\text{H}$  and  $^{133}\text{Cs}$  and a measurement comparing  $^{87}\text{Rb}$  and  $^1\text{H}$ , resulting in a stringent limit on LPI,  $\beta_{\text{Rb}} - \beta_{\text{H}} = (-2.7 \pm 4.9) \times 10^{-7}$ . The method of making these measurements for continuous clocks is discussed.

DOI: [10.1103/PhysRevA.87.010102](https://doi.org/10.1103/PhysRevA.87.010102)

PACS number(s): 11.30.Er, 04.80.Cc, 06.30.Ft

### I. INTRODUCTION

Metric theories of gravity are based on the universal coupling of gravity to matter and energy, formalized in the Einstein equivalence principle (EEP). This principle states that gravitational acceleration is independent of composition (the weak equivalence principle) and that nongravitational measurements should be independent of the velocity of the freely falling reference frame (local Lorentz invariance) and of the location in spacetime (local position invariance; LPI) where they are carried out. Most efforts to unify general relativity, the widely accepted metric theory of gravity, with quantum field theory involve a breakdown of EEP [1,2].

According to conventional parametrizations, the principle of LPI is the least well tested tenet of EEP [3]. Furthermore, tests of LPI can be used to constrain the coupling of certain fundamental constants to gravitational potential and, more generally, to test for spatial dependence of the constants. A spatial dependence is allowed under a particular scenario in string theory [2], and there have been published claims of a measured spatial variation of the fine structure constant from astronomical observations [4].

The traditional redshift for a clock with frequency  $\nu$  depends only on the change in gravitational potential  $\Delta U$ ,  $\Delta\nu/\nu = \Delta U/c^2$ , and should be independent of the clock composition, for example. Violation of LPI may manifest itself in an anomalous gravitational redshift of an atomic clock, commonly parameterized with a clock-dependent term  $\beta$ :

$$\Delta\nu/\nu = (1 + \beta)\Delta U/c^2. \quad (1)$$

Measurements of gravitational redshift vs  $\Delta U$  provide a test of the redshift formula and of LPI [5]. Alternatively, the relative redshift of two clocks of different composition can be used; if the two clocks have LPI-violating parameters  $\beta_1$  and  $\beta_2$ , the relative frequencies should vary with the gravitational potential as

$$(\Delta\nu/\nu)_{1,2} = (\beta_1 - \beta_2)\Delta U/c^2. \quad (2)$$

This differential measurement can be made with a higher precision than a direct measurement of redshift vs  $\Delta U$  at the possible expense of a reduction in the size of the effect if  $\beta_1$  and  $\beta_2$  are close in value. The gain in precision reported here compared to direct redshift measurements is about 200, indicating that the two redshift anomalies must be the same to  $\sim 5 \times 10^{-3}$  for the best LPI test by absolute measurement of the gravitational redshift to be comparable [5]. The possible range of values for  $\beta_1 - \beta_2$  is one reason why tests with many different pairs of clocks are valuable.

A convenient gravitational potential to use for an LPI test is the solar potential experienced on earth; due to the earth's elliptical orbit, this varies annually as

$$\Delta U_s/c^2 = A \sin(\omega t + \phi_0), \quad (3)$$

where  $A = 1.65 \times 10^{-10}$ ,  $\omega = 0.0172$  rad/day, and  $\phi_0$  is the phase such that  $\Delta U_s$  is a minimum at aphelion, such as modified Julian day 55746 (July 4, 2011). (The additional variation in potential due to the earth's rotation about its axis is  $\sim 10^{-3} \Delta U_s$  for our latitude and is ignored here.) The difference in LPI-violating parameters for two atomic clocks of different composition can be measured by determining the size of the annual oscillation with the appropriate phase in the clocks' relative frequency.

Over the past decade measurements of this sort have been made using a variety of pairs of atomic species [6–10], summarized in Fig. 1. In almost all previous measurements absolute frequencies of a certain type of atomic clock are compared to a cesium standard over many years, with typically several (fractional) frequency comparisons at the  $\sim 10^{-15}$  level per year. We implement a different approach, where each of the atomic clocks used is in continuous operation, providing a much higher rate of frequency comparisons and resulting in a more efficient LPI test. We measure the output frequencies of clocks based on hyperfine transitions in  $^{87}\text{Rb}$ ,  $^{133}\text{Cs}$ , and  $^1\text{H}$  over 1.5 years. Frequency comparisons between any two of the atomic species can be made with a statistical uncertainty of  $\sim 10^{-15}$  in a day or two, resulting in several hundred data points per year of precision comparable to previous tests.

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Report Documentation Page			Form Approved OMB No. 0704-0188		
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1. REPORT DATE <b>22 JAN 2013</b>		2. REPORT TYPE		3. DATES COVERED <b>00-00-2013 to 00-00-2013</b>	
4. TITLE AND SUBTITLE <b>Tests of local position invariance using continuously running atomic clocks</b>			5a. CONTRACT NUMBER		
			5b. GRANT NUMBER		
			5c. PROGRAM ELEMENT NUMBER		
6. AUTHOR(S)			5d. PROJECT NUMBER		
			5e. TASK NUMBER		
			5f. WORK UNIT NUMBER		
7. PERFORMING ORGANIZATION NAME(S) AND ADDRESS(ES) <b>United States Naval Observatory, Washington, DC, 20392</b>			8. PERFORMING ORGANIZATION REPORT NUMBER		
9. SPONSORING/MONITORING AGENCY NAME(S) AND ADDRESS(ES)			10. SPONSOR/MONITOR'S ACRONYM(S)		
			11. SPONSOR/MONITOR'S REPORT NUMBER(S)		
12. DISTRIBUTION/AVAILABILITY STATEMENT <b>Approved for public release; distribution unlimited</b>					
13. SUPPLEMENTARY NOTES <b>PHYSICAL REVIEW A 87, 010102(R) (2013)</b>					
14. ABSTRACT <b>Tests of local position invariance (LPI) made by comparing the relative redshift of atomic clocks based on different atoms have been carried out for a variety of pairs of atomic species. In most cases, several absolute frequency measurements per year are used to look for an annual signal, resulting in tests that can span on order of a decade. By using the output of continuously running clocks, we carry out LPI tests with comparable or higher precision after less than 1.5 years. These include new measurements of the difference in redshift anomalies <math>\Delta \nu_{\text{H}}</math> for hyperfine transitions in <math>^{87}\text{Rb}</math> and <math>^{133}\text{Cs}</math> and in <math>^1\text{H}</math> and <math>^{133}\text{Cs}</math> and a measurement comparing <math>^{87}\text{Rb}</math> and <math>^1\text{H}</math>, resulting in a stringent limit on LPI, <math>\Delta \nu_{\text{H}}/\nu_{\text{H}} = (2.7 \pm 4.9) \times 10^{-8}</math>. The method of making these measurements for continuous clocks is discussed.</b>					
15. SUBJECT TERMS					
16. SECURITY CLASSIFICATION OF:			17. LIMITATION OF ABSTRACT <b>Same as Report (SAR)</b>	18. NUMBER OF PAGES <b>4</b>	19a. NAME OF RESPONSIBLE PERSON
a. REPORT <b>unclassified</b>	b. ABSTRACT <b>unclassified</b>	c. THIS PAGE <b>unclassified</b>			

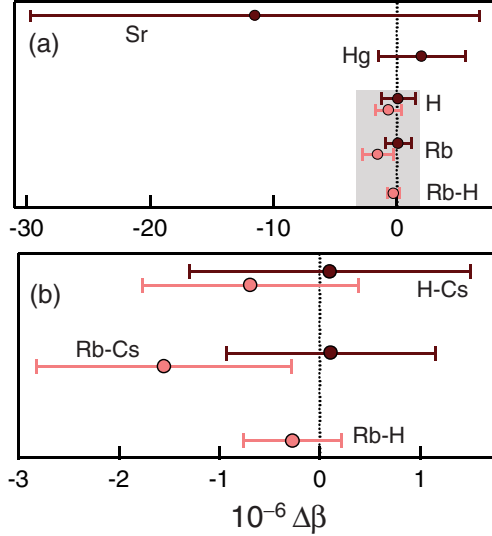


FIG. 1. (Color online) LPI tests for various pairs of atomic species in terms of the difference in anomalous redshift parameters,  $\Delta\beta = \beta_1 - \beta_2$ . (a) Dark data points are previous measurements: (i) neutral strontium optical transition against cesium standard, carried out over 2.5 years [9,11]; (ii) mercury ion optical transition against cesium standard, carried out over 5 years [7]; (iii) hydrogen hyperfine transition against cesium standard, carried out over 7 years [8]; and (iv) rubidium-87 hyperfine transition against cesium standard, carried out over 14 years [10]. Light data points are results presented in this work. (b) A more detailed look at the shaded region in (a).

## II. CLOCKS, METHOD, AND ANALYSIS

Because we do not evaluate absolute frequencies, we have to be confident that our clocks' outputs are stable or change in knowable, predictable ways. This is accomplished by maintaining a continuous record of a clock's frequency against a variety of other clocks, enabling us to determine if discontinuous changes in frequency or drift rate occur.

The US Naval Observatory (USNO) has a large ensemble of atomic clocks used for precise time [12], including commercial cesium-beam clocks, hydrogen masers, and recently added rubidium-fountain clocks, which have been in operation for 1.5 years [13]. All clocks are measured regularly against a master reference, which enables a comparison between any two clocks. The collection of each type of clock can be used to generate one or more stationary outputs representing that ensemble.

The rubidium frequency record used for LPI tests is an average of the frequencies of the two highest performing fountains over the past 1.5 years. Comparisons among rubidium fountains demonstrate that these two clocks exhibit extremely stable frequencies, with white-frequency noise levels below  $2 \times 10^{-13}$ . The fountains show no indication of frequency drift at the level of  $3 \times 10^{-18}$ /day, as determined by using cesium-fountain primary standards that contribute to International Atomic Time as a reference. As a baseline, we look for an LPI-violating signal in the frequency comparison of the two rubidium fountains. The result is 0 well within the error bar, indicating no bias in measurements using these clocks.

About 70 commercial cesium clocks are intercompared to create a single output used for our LPI analysis [14].

The output of this cesium ensemble is characterized by a white-frequency noise level of  $1 \times 10^{-12}$  and shows a drift in fractional frequency of  $2 \times 10^{-17}$ /day over the period in question. There is no long frequency record for an individual cesium clock without adjustments based on the performance of the ensemble, so there is no measure of an LPI-violating baseline for the individual clocks.

Hydrogen masers generally exhibit complicated behavior over long averaging times (months). They are known to be sensitive to environmental factors, which can cause frequency discontinuities in addition to less violent perturbations, and they exhibit significant frequency drifts, which may not be linear and which can change in time. Of the  $\sim 25$  hydrogen masers available to us, a subset is ruled out for these LPI tests because they exhibit discontinuous changes in frequency or drift rate during the interval of time that we are interested in. Further assessing which masers to include requires careful analysis of the frequency records of different masers compared against each other as well as against the other clocks used in the analysis, as discussed further below. The maser short-term frequency stability is better than that of either of the other types of clock used for these LPI tests. The white-frequency noise level of a maser-fountain comparison is typically of the order of  $5 \times 10^{-13}$ , limited by noise in the frequency-counter measurement system. Some masers are also measured with a dual-mixer measurement system, in which case a maser-fountain comparison exhibits a white-frequency noise level limited by fountain performance.

In order to measure  $\beta_1 - \beta_2$  for a pair of atomic species we need to examine the relative frequency record for an annual oscillation with phase  $\phi_0$  in the presence of a drift. We limit our analysis to clocks that show a frequency drift that is linear, which, for a  $\sim 1.5$ -year interval, has little impact on the determination of the sinusoid amplitude. Applying a least-squares fit using a function that is a combination of drift and oscillation gives results identical to fitting the frequency record to a line and then fitting those residuals to a sinusoid. We use the latter procedure for presenting data in the figures below. All relative frequency measurements participating in the fit derive from averaging intervals resulting in white-frequency noise, consistent with application of least-squares curve fitting.

### A. Rubidium vs cesium

Application of this procedure to a rubidium-cesium LPI test is straightforward. Figure 2(a) shows the fractional frequency difference between the cesium ensemble and the average of rubidium fountains over 1.5 years. Raw data are recorded at hourly intervals; the plots in the figure show the frequency averaged over 2 days. The graph includes a fit indicating a relative drift of  $2 \times 10^{-17}$ /day. In Fig. 2(b), the residuals after removing the drift are fit to an LPI-violating signal, shown amplified in the graph. The amplitude for this fit is  $\Delta\nu_{\text{Rb}}/\nu_{\text{Rb}} - \Delta\nu_{\text{Cs}}/\nu_{\text{Cs}} = (2.6 \pm 2.1) \times 10^{-16}$  (all uncertainties are 1 standard deviation). This translates to a difference in redshift anomalies of

$$\beta_{\text{Rb}} - \beta_{\text{Cs}} = (-1.6 \pm 1.3) \times 10^{-6}. \quad (4)$$

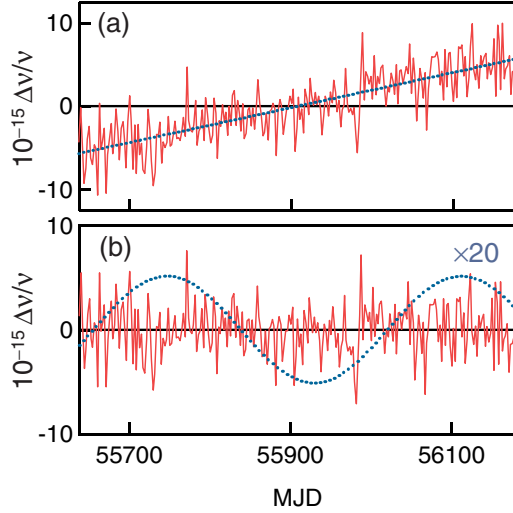


FIG. 2. (Color online) (a) Frequency difference of rubidium-fountain average and cesium ensemble over 1.5 years. Fit to a linear drift is shown. (b) Residuals after removing drift, along with the fit to an LPI-violating signal, amplified  $\times 20$ . MJD stands for modified Julian day.

### B. Hydrogen

From the subset of masers that exhibit no apparent discontinuities in frequency or drift over all or a large portion of the past 1.5 years, we need to establish which masers have frequency records that are truly linear, without higher order terms. Some masers obviously exhibit frequencies that change in time in a nonlinear way and can be removed from consideration. To rule out masers with a more subtle nonlinear behavior, we rely on the quality of the least-squares fit to the relative frequency record of a maser and the rubidium average [15]. If the data do not fit well, as indicated by the reduced  $\chi^2$  and a visual inspection of the residuals, the maser is not included in the analysis. We attribute a poor fit to unmodeled behavior in the maser; this procedure leaves us with four masers in the LPI analysis.

It is possible that these four masers still display unmodeled behavior, which mimics an LPI-violating signal. Such unmodeled behavior would not be surprising since the phase of the LPI-violating signal is close to that of the seasonal cycle. We assess this by looking at all six possible pairs of the four masers for an LPI-violating signal. These measurements enable a  $\beta$  to be determined for each maser that serves as a baseline for remaining unmodeled behavior. Each  $\beta$  value serves as a bias and is subtracted from the measured LPI-violating signal for each maser-rubidium or maser-cesium comparison. Once this bias is included in a clock comparison using a given maser, the final uncertainty for that comparison is the combined statistical uncertainty from the LPI fit and the uncertainty associated with the bias.

In order to account for the different uncertainties associated with each maser, the hydrogen results are determined by measuring LPI-violating amplitudes for each of the four masers against the cesium and rubidium clocks and combining them in a weighted average [16]. The four measurements against the cesium ensemble give

$$\beta_H - \beta_{Cs} = (-0.7 \pm 1.1) \times 10^{-6}. \quad (5)$$

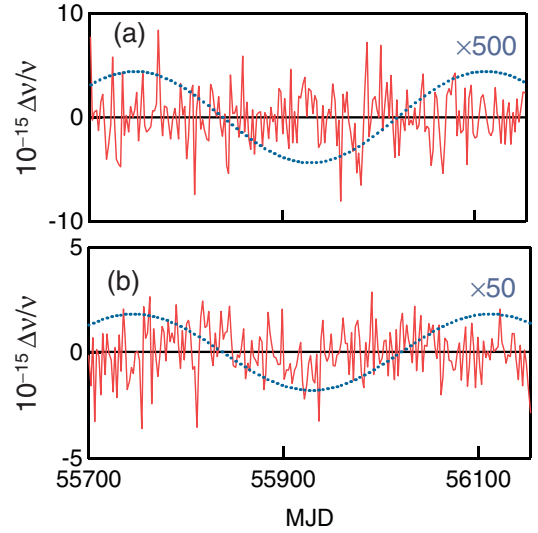


FIG. 3. (Color online) Frequency comparison for one hydrogen maser and (a) the cesium ensemble and (b) the rubidium-fountain average. Each point is a 2-day average, and linear drifts have been removed. Fits to LPI-violating signals are amplified by the factors shown.

The average of the rubidium fountains measured against the four masers yields

$$\beta_{Rb} - \beta_H = (-2.7 \pm 4.9) \times 10^{-7}. \quad (6)$$

Figure 3 shows the frequency of one of the masers compared against [Fig. 3(a)] the cesium ensemble and [Fig. 3(b)] the rubidium average.

### III. DISCUSSION AND SUMMARY

Measurements of  $\beta_1 - \beta_2$  for different pairs of atomic species can be used to constrain the coupling to gravitational potential of certain dimensionless constants: the fine structure constant  $\alpha$ , the electron-to-proton mass ratio  $m_e/m_p$ , and the ratio of the light quark mass to the quantum chromodynamics length scale,  $m_q/\Lambda_{QCD}$ , where  $m_q$  is the average of the up and down quark masses [17]. For the case of massive bodies the size of the sun, the coupling of these constants to gravity is expected to be proportional to the coupling to a hypothetical scalar field that is a component of many cosmological models and that could affect their values [2,17].

The coupling of these fundamental constants to gravitational potential can be characterized by dimensionless coupling parameters  $k_\epsilon$  [17]:

$$\frac{\delta\epsilon}{\epsilon} = k_\epsilon \left( \frac{\Delta U}{c^2} \right), \quad (7)$$

where  $\epsilon$  stands for the three constants  $\alpha$ ,  $m_e/m_p$ , and  $m_q/\Lambda_{QCD}$ . Using calculated sensitivities of different atomic transition frequencies to variations in these constants, differences in redshift anomalies for a pair of atomic species can be related to linear combinations of the coupling parameters  $k_\epsilon$  [10,18,19]. The three LPI measurements presented here can be used to put limits on the coupling of  $\alpha$  and  $m_q/\Lambda_{QCD}$  using only clocks from one institution,  $k_\alpha = (3.3 \pm 2.3) \times 10^{-6}$  and  $k_{m_q/\Lambda_{QCD}} = (-1.8 \pm 1.3) \times 10^{-5}$ . Measurements between

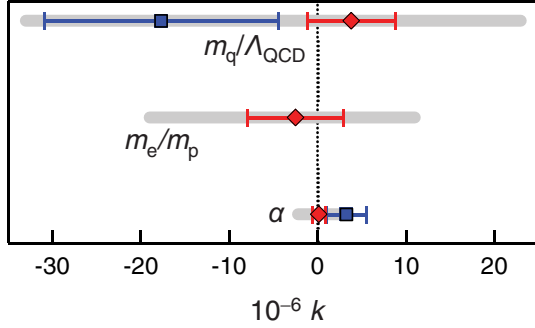


FIG. 4. (Color online) Constraints on coupling of  $\alpha$ ,  $m_e/m_p$ , and  $m_q/\Lambda_{\text{QCD}}$  to gravitational potential characterized by the dimensionless parameters  $k_\alpha$ ,  $k_{m_e/m_p}$ , and  $k_{m_q/\Lambda_{\text{QCD}}}$ . Gray bands show constraints imposed by previous measurements, compiled from Refs. [7–10] and presented in [10]. Diamonds show the tighter constraints that result when our measurements are included in the analysis. Squares represent the values for  $k_\alpha$  and  $k_{m_q/\Lambda_{\text{QCD}}}$  that are derived using only the results presented here.

optical and microwave clocks are needed to constrain the coupling for  $m_e/m_p$ . By using our measurements along with those from Refs. [7–10] we get limits on all three couplings:

$$\begin{aligned} k_\alpha &= (1.7 \pm 7.5) \times 10^{-7}, \\ k_{m_e/m_p} &= (-2.5 \pm 5.4) \times 10^{-6}, \\ k_{m_q/\Lambda_{\text{QCD}}} &= (3.8 \pm 4.9) \times 10^{-6}. \end{aligned} \quad (8)$$

These results are plotted in Fig. 4.

The tests involving cesium are limited by the white-frequency noise of the cesium ensemble, so the precision of these measurements should improve with longer fitting times. The limit on the precision of the rubidium-hydrogen test is a combination of white-frequency noise from the measurement

system and uncertainty in the LPI-mimicking amplitude for each maser. While these uncertainties should also improve with longer fitting times, the unpredictable nature of masers makes this less certain. Even now, with an uncertainty on the rubidium-hydrogen measurement in the mid  $10^{-7}$ s, this very precise LPI test could remain competitive for the near future. While relative frequency measurements of two optical clocks should yield much more stringent limits, many optical clock frequencies are measured against the cesium standard, limiting the precision to a level governed by a microwave transition. In terms of coupling parameters, optical-to-optical measurements only constrain  $k_\alpha$  [10]; measurements involving hyperfine transitions are required to determine  $k_{m_e/m_p}$  and  $k_{m_q/\Lambda_{\text{QCD}}}$ . Furthermore, our continuous clock measurements are competitive with LPI tests that could be implemented in near-term space missions. Clock experiments intended for the International Space Station [20] could improve upon the precision of absolute redshift measurements but not differential measurements; the solar-potential driving term is the same as on earth, and there is no terrestrial driving term due to the station's circular orbit. In terms of space clocks, improvements to differential redshift tests of LPI could result from proposed missions that would put clocks in a highly eccentric earth orbit [21] or a solar system escape trajectory [22].

In summary, we have made precise LPI tests with three different pairs of atomic species using continuously running atomic clocks and have used these to tighten the constraints on coupling of fundamental constants to gravity.

## ACKNOWLEDGMENTS

We are indebted to Paul Koppang for providing clock data and an understanding of the cesium ensemble, and we benefited from discussions with Demetrios Matsakis. Atomic fountain development at USNO has been funded by ONR and SPAWAR.

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- [1] T. Damour, *Class. Quantum Grav.* **13**, A33 (1996).
  - [2] T. Damour and A. M. Polyakov, *Nucl. Phys. B* **423**, 532 (1994).
  - [3] C. M. Will, *Living Rev. Relativity* **9**, 3 (2006).
  - [4] J. K. Webb, J. A. King, M. T. Murphy, V. V. Flambaum, R. F. Carswell, and M. B. Bainbridge, *Phys. Rev. Lett.* **107**, 191101 (2011).
  - [5] R. Vessot *et al.*, *Phys. Rev. Lett.* **45**, 2081 (1980).
  - [6] S. J. Ferrell *et al.*, *Phys. Rev. A* **76**, 062104 (2007).
  - [7] T. M. Fortier *et al.*, *Phys. Rev. Lett.* **98**, 070801 (2007).
  - [8] N. Ashby, T. P. Heavner, S. R. Jefferts, T. E. Parker, A. G. Radnaev, and Y. O. Dudin, *Phys. Rev. Lett.* **98**, 070802 (2007).
  - [9] S. Blatt *et al.*, *Phys. Rev. Lett.* **100**, 140801 (2008).
  - [10] J. Guéna *et al.*, *Phys. Rev. Lett.* **109**, 080801 (2012).
  - [11] The strontium uncertainty shown is a factor of 2 larger than published in [9]. See also [10].
  - [12] D. Matsakis, in *Proceedings of the Frequency Control Symposium, 2009 Joint with the 22nd European Frequency and Time Forum* (IEEE International, New York, 2009), pp. 653–658.
  - [13] S. Peil, S. Crane, J. Hanssen, T. Swanson, and C. Ekstrom, in *Proceedings of the 2012 IEEE International Frequency Control Symposium* (IEEE International, New York, 2012), p. 1.
  - [14] Using the commercial cesium clocks at USNO is more straightforward than using the cesium fountains that contribute to International Atomic Time, although measurements against the world's primary standards may be considered in future work.
  - [15] We rely on the comparison against a fountain because it is lower noise than the comparison against the cesium average.
  - [16] For rubidium, individual maser measurements are uncorrelated. For cesium, correlation between individual measurements gets taken into account.
  - [17] V. Flambaum, *Int. J. Mod. Phys. A* **22**, 4937 (2007).
  - [18] V. V. Flambaum and A. F. Tedesco, *Phys. Rev. C* **73**, 055501 (2006).
  - [19] T. H. Dinh, A. Dunning, V. A. Dzuba, and V. V. Flambaum, *Phys. Rev. A* **79**, 054102 (2009).
  - [20] L. Cacciapuoti *et al.*, *Nucl. Phys. B: Proc. Suppl.* **166**, 303 (2007).
  - [21] S. Schiller *et al.*, *Exp. Astron.* **23**, 573 (2008).
  - [22] P. Wolf *et al.*, *Exp. Astron.* **23**, 651 (2009).